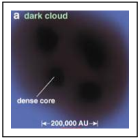
Disks in Astrophysics

Brian Kloppenborg

Draft: October 15, 2009

Outline

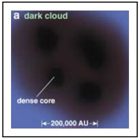
Single Star Formation

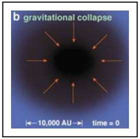


Images Courtsey of SSC IR Compendium

1. Cloud of gas and dust

Single Star Formation

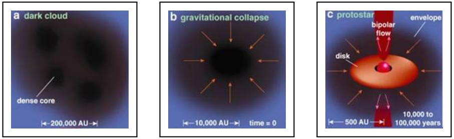




Images Courtsey of SSC IR Compendium

1. Cloud of gas and dust
2. Gravitational collapse

Single Star Formation



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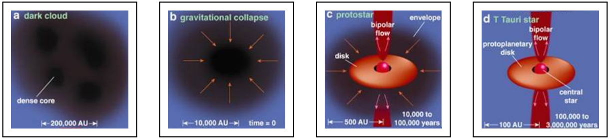
1. Cloud of gas and dust
2. Gravitational collapse
3. Conservation of angular momentum and collisions cause disk to form.

Single Star Formation



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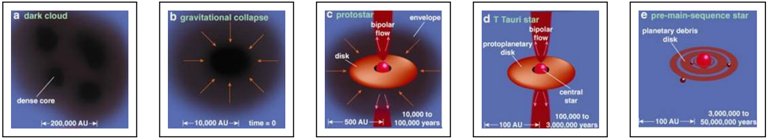
1. Cloud of gas and dust
2. Gravitational collapse
3. Conservation of angular momentum and collisions cause disk to form.
4. Envelope has dissapated or collapsed into the disk.

Single Star Formation



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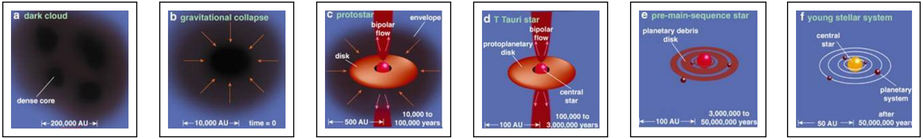
Single Star Formation



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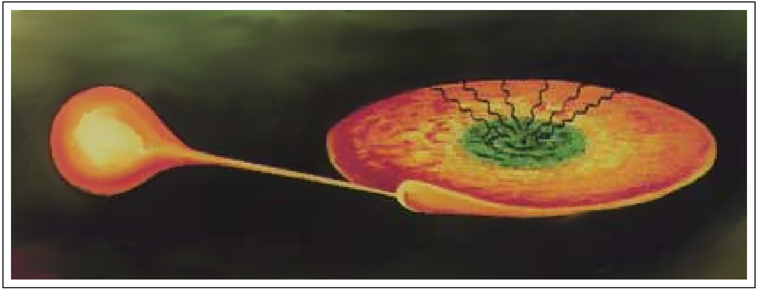
1. Cloud of gas and dust
2. Gravitational collapse
3. Conservation of angular momentum and collisions cause disk to form.
4. Envelope has dissapated or collapsed into the disk.
5. Collisions inside disk cause planetesimals for form, clearing the disk of debris.
6. Star ignites hydrogen in its core.

Binary Star Evolution



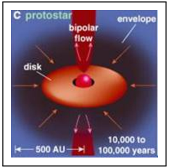
Artists impression of 4U 1820-30, image courtsey of NASAs HEASARC

Like single star evolution except:

- Roche Lobe Overflowing
- Mass Transfer Streams
- Ensuing Disk Hot Spots

Disk Evolution

Artist Impressions of Disk Evolution



Images Courtesy of STScI

Disk Evolution

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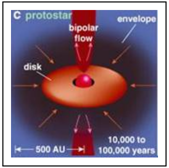

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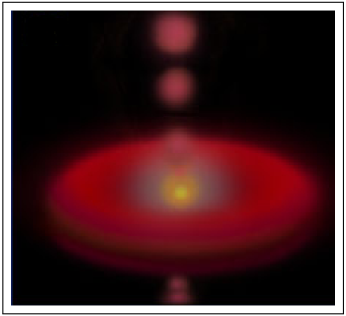

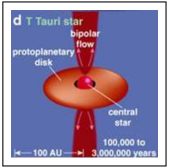

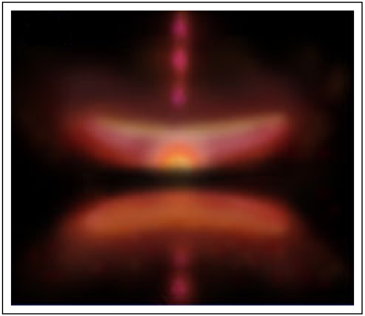

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Disk Evolution

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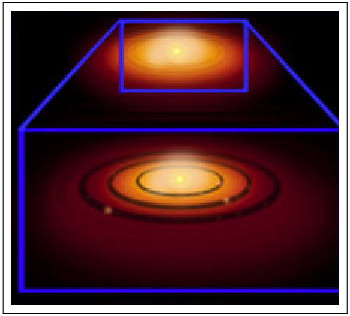
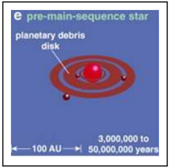
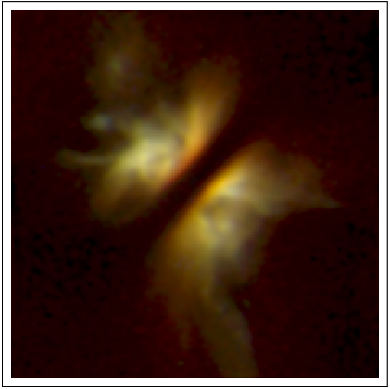




Images Courtesy of STScI

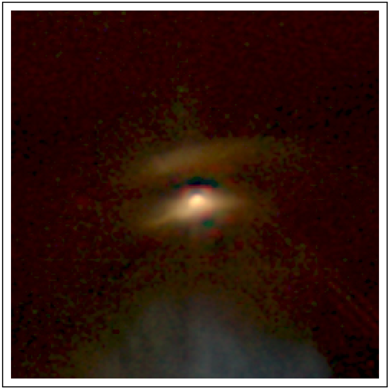
Direct Imaging

Butterfly Star



Karl Stapelfeldt (JPL) and colleagues, and NASA

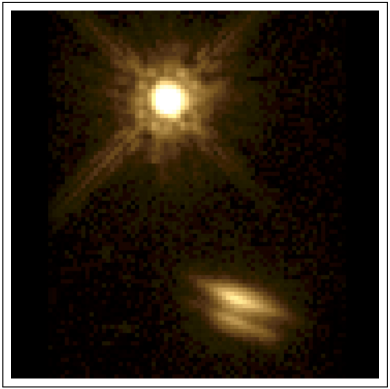
Herbig-Haro (HH) 6-5



D. Padgett (IPAC/Caltech), W. Brandner (IPAC), K. Stapelfeldt (JPL) and NASA

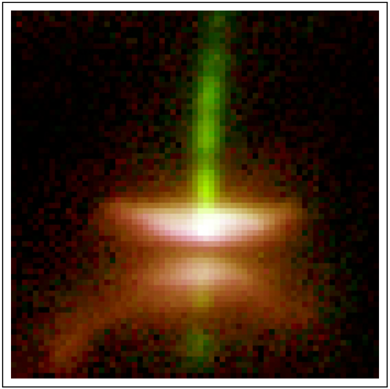
Direct Imaging

HK Tauri



Karl Stapelfeldt (JPL) and colleagues, and NASA

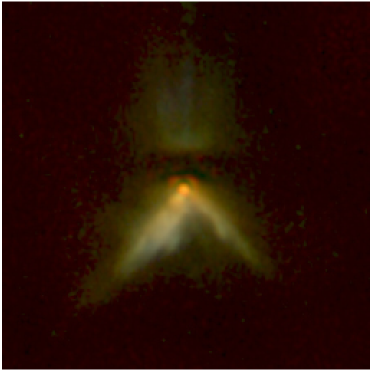
HH 30



Chris Burrows (STScI), the WFPC2 Science Team and NASA

Direct Imaging

DG Tauri B



(IR), D. Padgett (IPAC/Caltech), W. Brandner (IPAC), K. Stapelfeldt (JPL) and NASA

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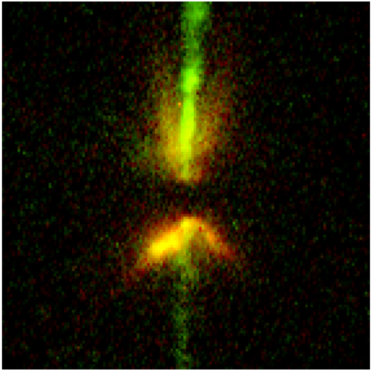

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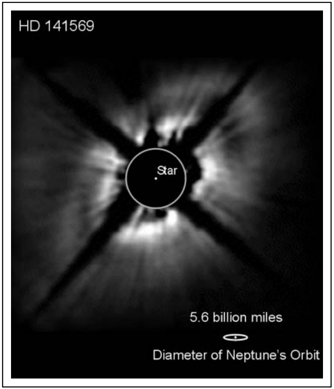
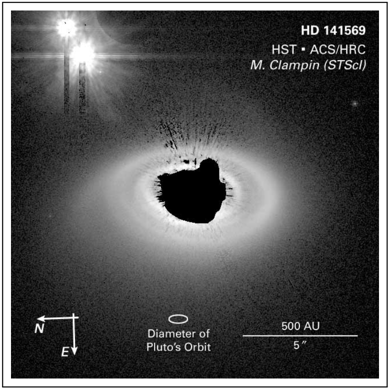
Direct Imaging

HD 141569



Near IR), B. Smith (U. Hawaii), A. Weinberger, E. Becklin (UCLA), and G. Schneider (U. Arizona)

HD 141569



(Hubble), M. Clampin (NASA Goddard)

Spectral Energy Distributions (SEDs)

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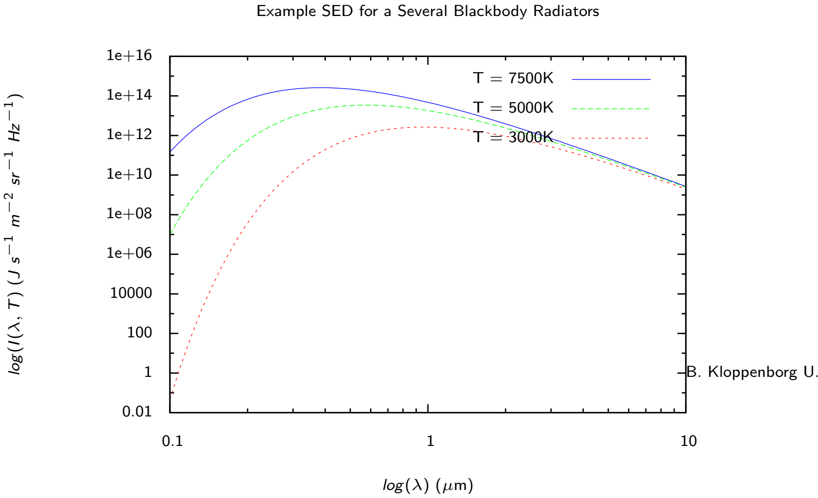
Blackbody Radiation is a First-Order approximation for radiation from objects. Blackbody radiation is described by Plancks Law:

Planck's Law

$$I(\lambda, T) = \frac{2hc^2}{\lambda^5} \frac{1}{e^{\frac{hc}{\lambda k T}} - 1} \tag{1}$$

Spectral Energy Distributions (SEDs)



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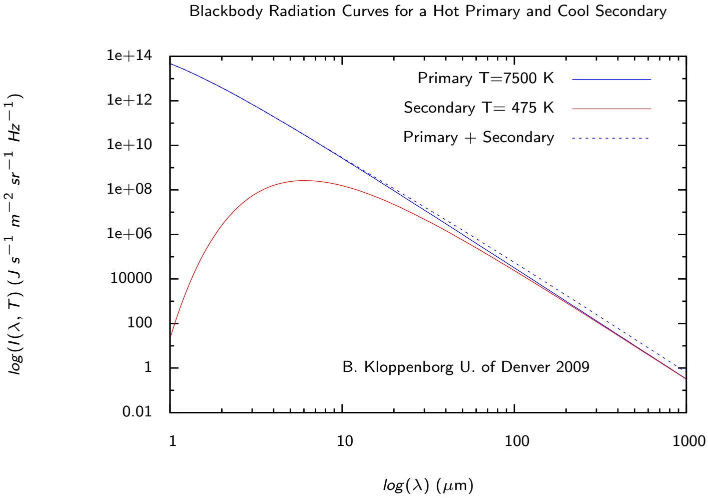
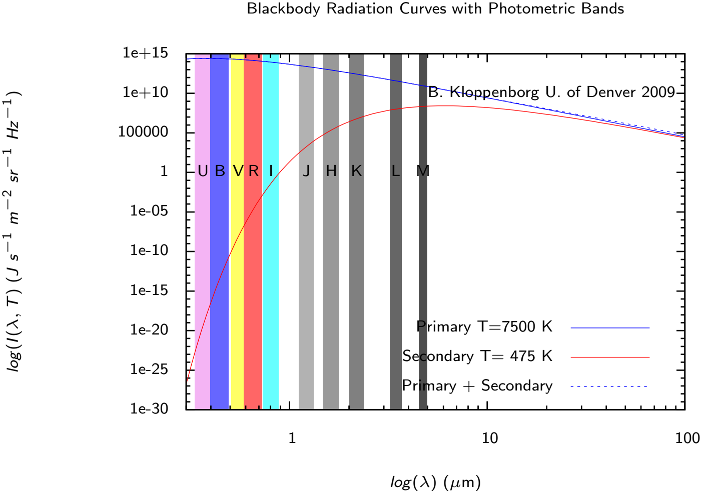


Photometery

Photometry measures the brightness of stars in specific passbands.

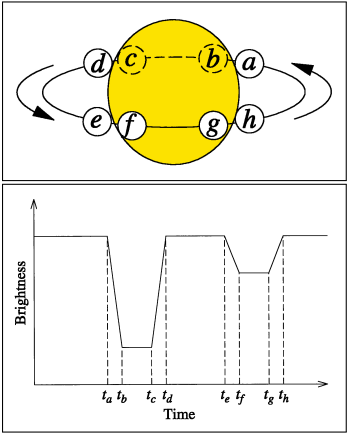
Photometery



Photometery

Photometery can Determine

- Period of variable stars,minor planets, AGNs,transiting extrasolar planets.



Eclipsing Binary With Light Curve (Ostlie 1996)

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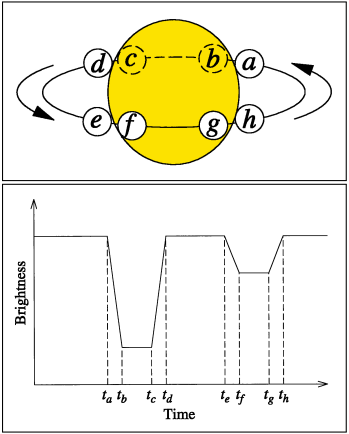

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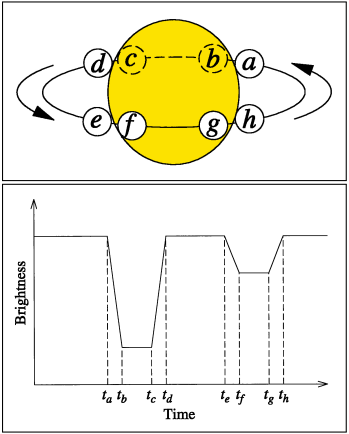

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- Period of variable stars,minor planets, AGNs,transiting extrasolar planets.
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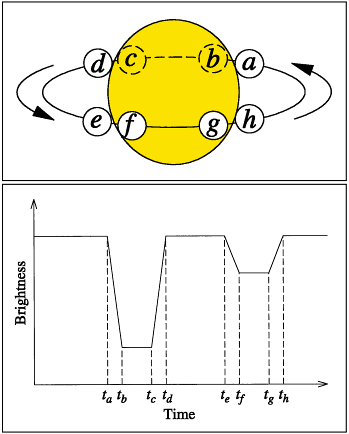
- Blackbody Temperature of an object.
- Total energy output of supernovae.



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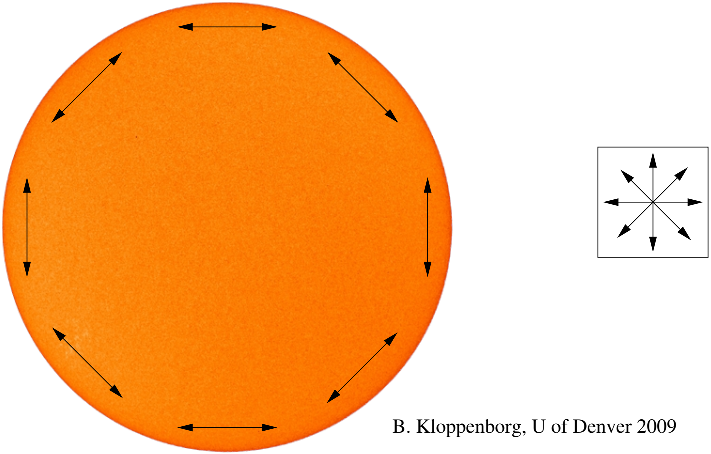
Polarimetery

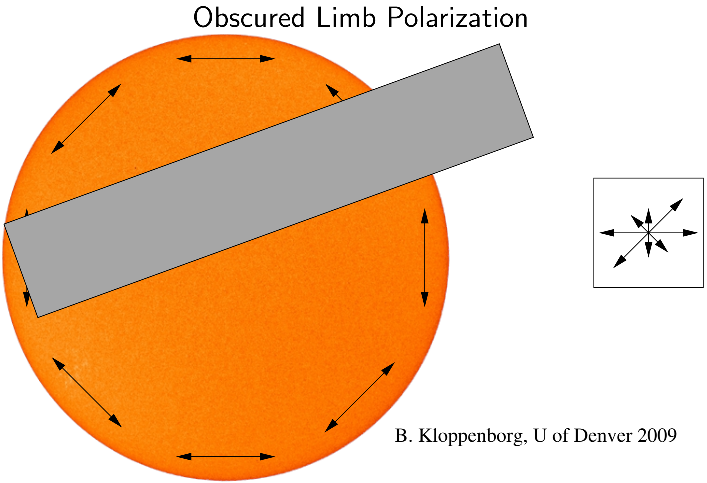
Polarimetery

Limb Polarization



Limb Polarization.
MDI Continuum image of the Sun courtesy of SOHO

Polarimetery



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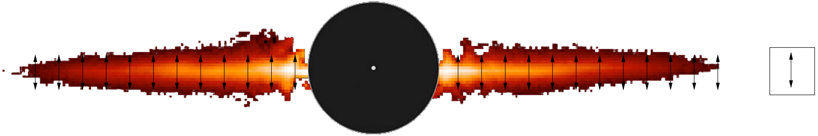
Polarimetery

A Disk Can Cause a Net Polarization



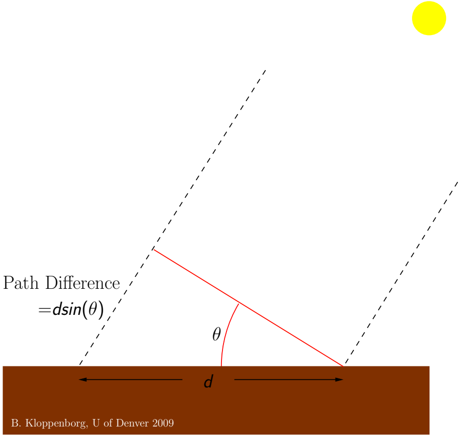
Polarization from a Disk.
HST Image of AU Mic Debris Disk courtesy of NASA, ESA, and J. Graham (UC, Berkeley)

Interferometery

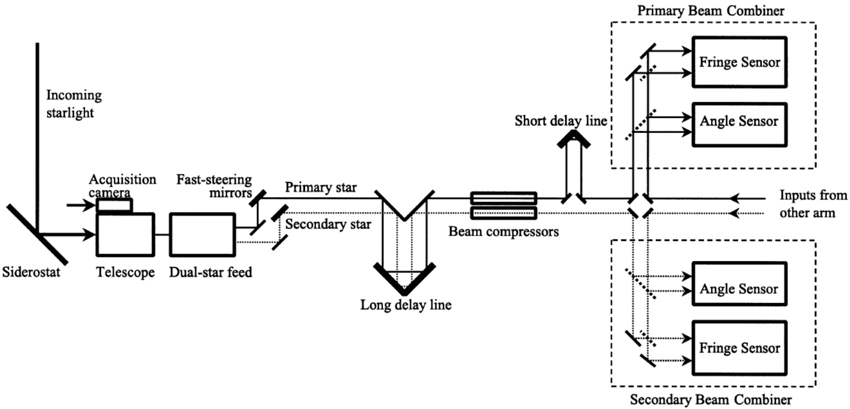
Interferometry provides a method of indirectly imaging objects too small for traditional imaging devices.

Interferometery

Interferometer Diagram

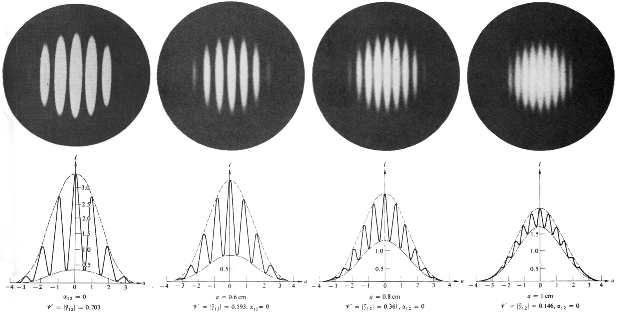


Interferometery



Schematic Drawing of PTI Optical Components (Colavita, 1999)

Fringes

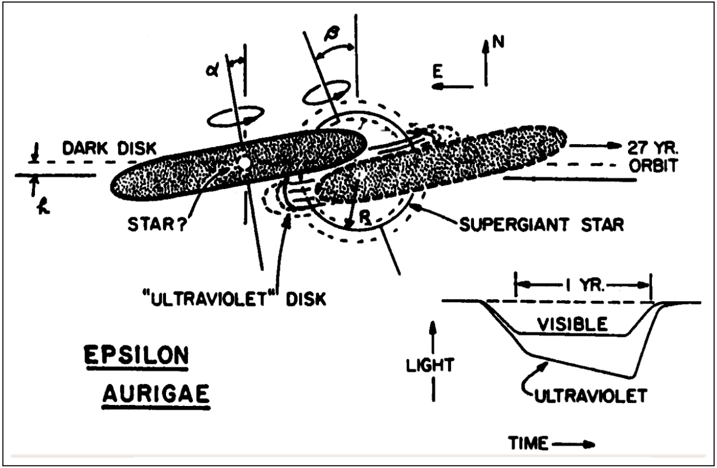


Fringes as seen by an Interferometer (Hecht, 2002)

Visibility Squared:

$$V^2 = \left(\frac{I_{max} - I_{min}}{I_{max} + I_{min}}\right)^2$$

Current Model of ϵ Aurigae



Model of ϵ Aurigae System (NASA, 1985)

Photometery Observations

Photometric Observations during Eclipse

1848 Schmidt (see Gussow, 1936), others?

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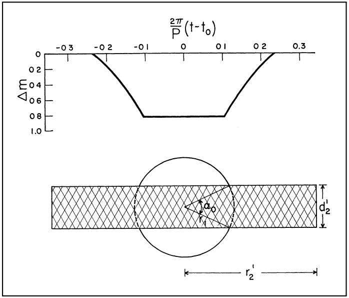
2009 *Your Name Here*

Models from Photometric Data

1965 Huang: Dark rectangular object moving across star face.



Schematic Diagram for Aur with light curve (Huang, 1965)

Models from Photometric Data

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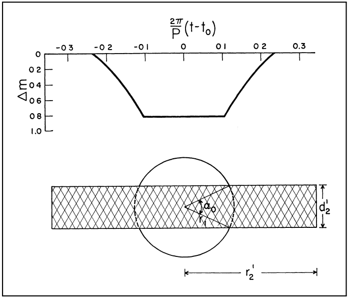

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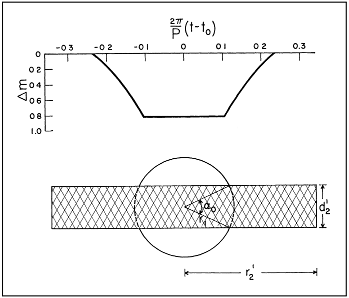
- 1974 Huang: Comparison of Thin/Thick Disk Models with Observed Light Curves.



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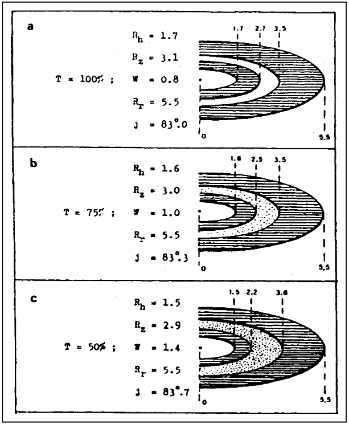
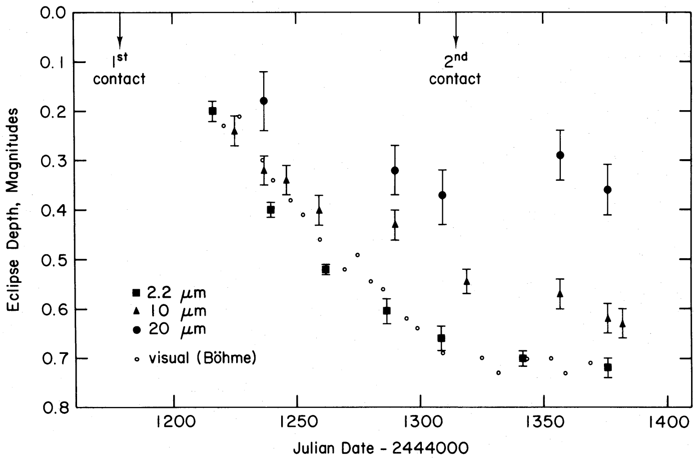
- 1990 Ferluga: Disk with Semitransparent Rings.



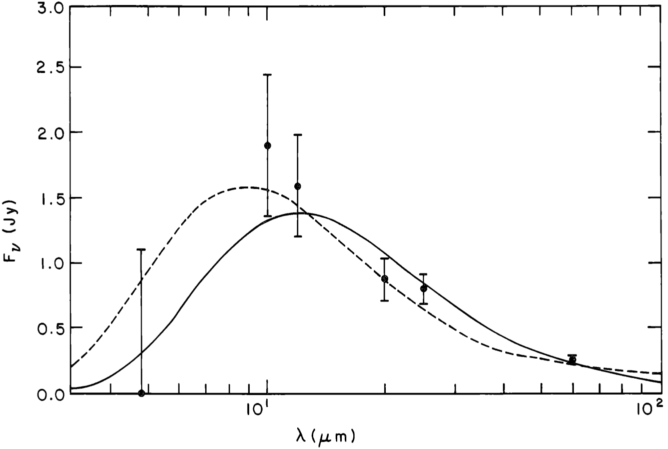
Three best-fitting ringed disk models to photometric curves. Units are in AU, j is the inclination angles (exaggerated in figures) (Ferluga, 1990)

Photometery



IR Excess as reported by Backman (1984)

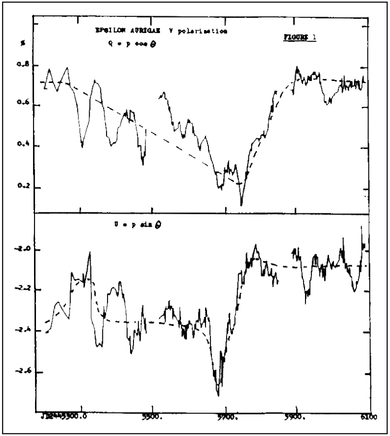
SEDs



IR Excecess as found by IRAS (Backman, 1985). *Solid Line* opaque secondary, $T = 475$K, $\Omega = 8.6 \times 10^{-16}$ sr; *Dashed Line* optically thin secondary $T = 575$K, $\Omega = 4.4 \times 10^{-16}$ sr, particle radius 5.1 μm.

Polarimetery

1985 Kemp: Polarization Curve During Eclipse with preliminary model.

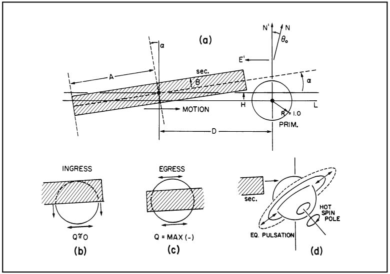


In-eclipse Polarization Data for Aur with light curve (Kemp, 1986)

Polarimetery

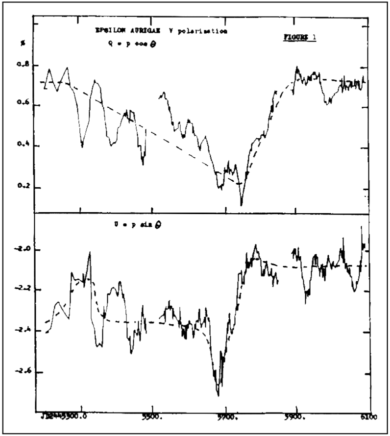
1985 Kemp: Polarization Curve During Eclipse with preliminary model.

1986 Kemp: Detailed Model for Polarization Curve.



Model for Polarization Data (Kemp, 1986)

Polarimetery



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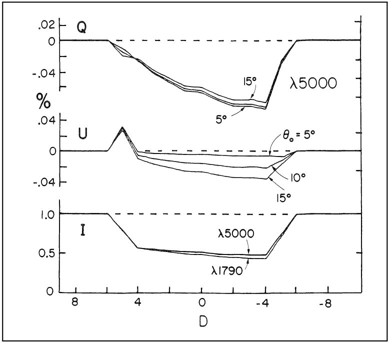


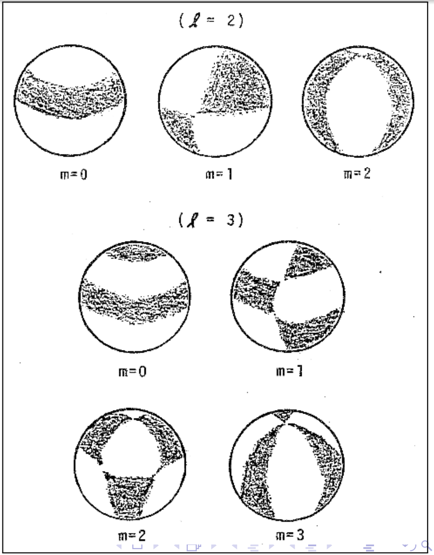
Theoretical Curves for Kemps Model. (Kemp, 1986)

Polarimetery

1985 Kemp: Polarization Curve During Eclipse with preliminary model.

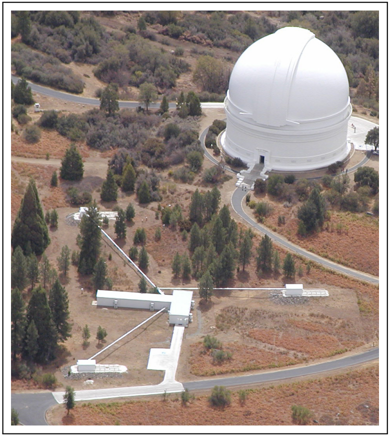
1986 Kemp: Detailed Model for Polarization Curve.

1989 Henson: Possible detection of non-radial pulsation.



Right: Possible Pulsation Modes of Aur (Henson, 1989)

Palomar Testbed Interferometer



Aerial View of PTI and the 200" Palomar Telescope (Gerald van Belle)

- PTI Operated by the Michelson Science Center on behalf of CalTech and NASA-JPL
- Maximum Baseline, 110 meters
- Resolution 1.67 - 2.18 mas (8.1 - 10.5 nano-radians)

PTI: Reduced Data

UTDate, JD-2,450,000	GMT start	Baseline*	Nscans	Mode	V^2	UDD [mas]	Error [mas]	V [mag]
2007Oct19, 4393	09:57	NS	14	K-low	0.516	2.19	0.06	3.036
2007Oct20, 4394	10:21	NS	6	K-high	0.544	2.16	0.12	3.036
2007Oct21, 4395	10:45	NS	3	K-low	0.583	1.90	0.13	3.036
2007Dec23, 4458	04:41	NW	6	K-low	0.574	2.36	0.14	3.046
2007Dec24, 4459	04:48	NW	6	K-low	0.565	2.37	0.11	3.043
2008Feb16, 4513	03:05	NW	2	K-low	0.527	2.60	0.15	2.98
2008Feb17, 4514	04:48	NW	5	K-low	0.572	2.28	0.15	2.98
2008Feb18, 4515	03:01	NW	5	K-low	0.624	2.25	0.12	2.98
New Data								
2008Oct26, 4765	08:20	NW	4	K-low	0.609	2.35	0.16	3.052
2008Oct26, 4765	08:30	NS	5	K-low	0.491	2.16	0.08	3.052
2008Nov8, 4778	08:49	NS	12	K-low	0.435	2.34	0.08	3.057
2009Nov9, 4779	09:22	NW	1	K-low	0.462	2.86	0.10	3.054
Archival Data								
1997Oct22, 0744	11:54	NS	1	K-low	0.376	2.50	0.17	2.986
1997Nov09, 0762	09:38	NS	2	K-low	0.438	2.32	0.09	2.977
1998Nov07, 1125	10:25	NS	4	K-low	0.515	2.09	0.10	2.997
1998Nov25, 1143	10:19	NS	2	K-low	0.458	2.25	0.08	2.998
2005Dec11, 3715	06:33	NW	86	Insufficient Data Points				3.02
2006Jan31, 3766	04:27	NW	86	No Cal Stars				3.08

Diameters obtained from Wide-Band Visibility mode data. *N-S baseline, 109 meters; N-W baseline, 86 meters.

Data prior to Oct. 2008 published in Stencel et. al 2008. V-band data courtsey of Jeffery Hopkins.

CHARA

Mt. Wilson Today, (Georgia State University)

- Operated by Georgia State University and collaborators.
- Six 1-meter Telescopes
- 15 possible baselines from 31 to 331 meters
- One of two operating ranges: 2.0 - 2.5 μm
- 0.6 mas resolution

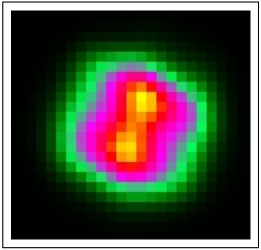
Observations and Results

Observations:

2008-09-19, 2008-11-07, 2008-11-08, 2008-12-10

Observations and Results

Observations:
2008-09-19, 2008-11-07, 2008-11-08, 2008-12-10

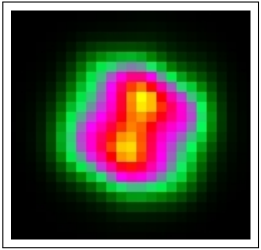


BSMEM

Scale: 0.16 $\frac{mas}{pixel}$
$\approx$ 1 nanoradian

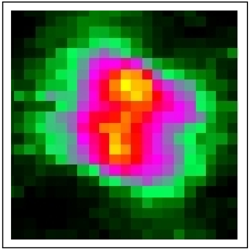
Observations and Results

Observations:
2008-09-19, 2008-11-07, 2008-11-08, 2008-12-10



BSMEM

Scale: 0.16 $\frac{mas}{pixel}$
$\approx$ 1 nanoradian



MACIM

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